

The background of the slide is a dark, deep-space photograph of a galaxy, likely NGC 4526, with a prominent bright supernova remnant (SN 1994D) visible as a glowing, irregularly shaped region. The galaxy's spiral arms are faintly visible against the dark sky.

β -decay measurements of r-process nuclei

Fernando Montes

Michigan State University

National Superconducting National Laboratory

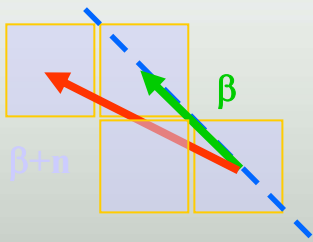
Joint Institute for Nuclear Astrophysics JINA

Supernova 1994D in NGC 4526

Nuclear physics in the r-process

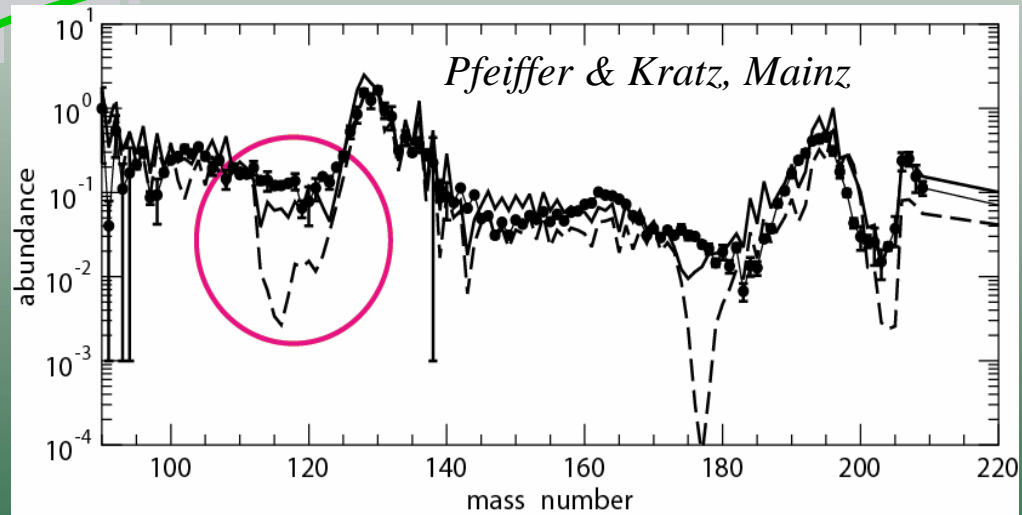
Fission rates and distributions

β -delayed n-emission branchings
(final abundances)

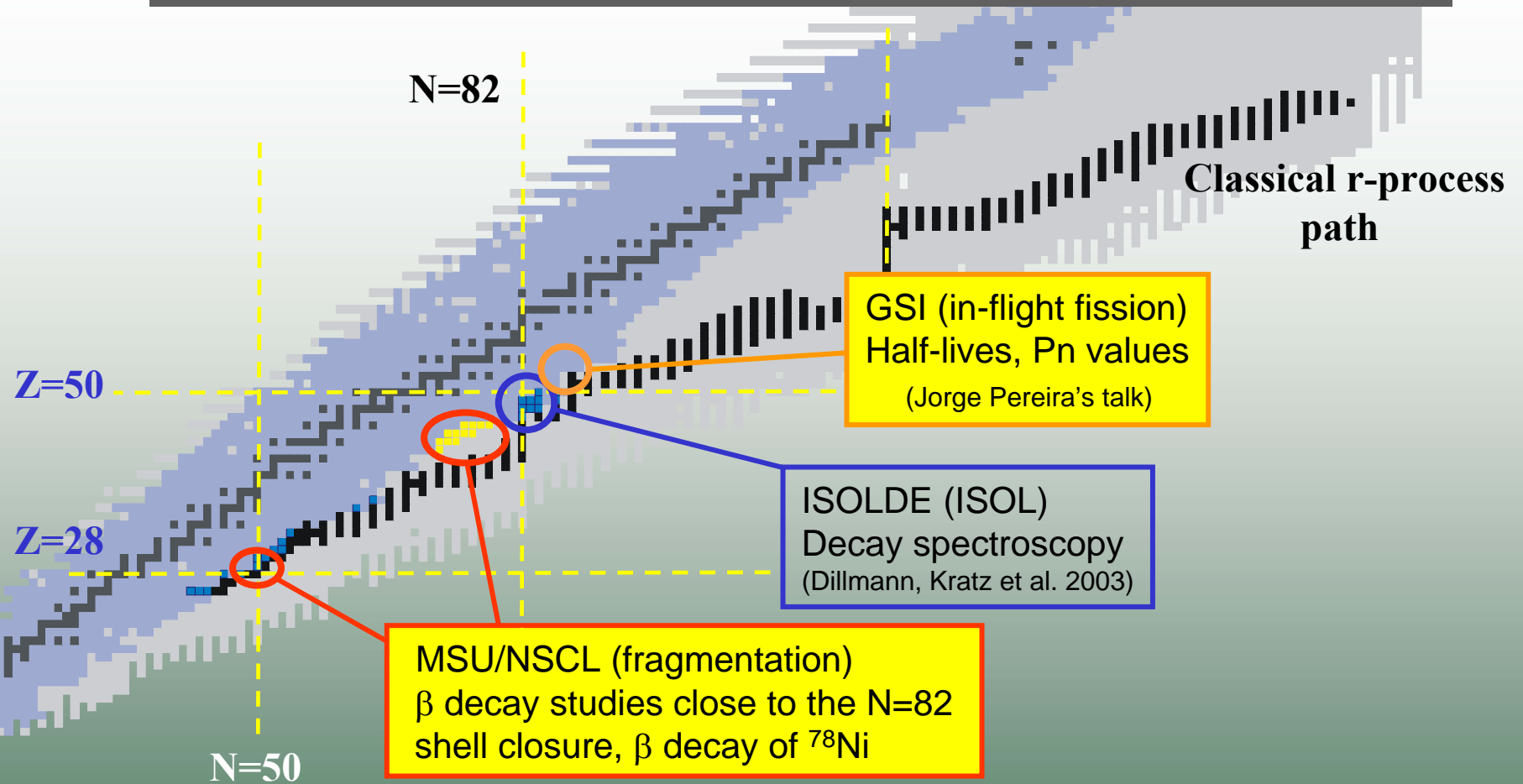


β -decay half-lives
(abundances and
process speed)

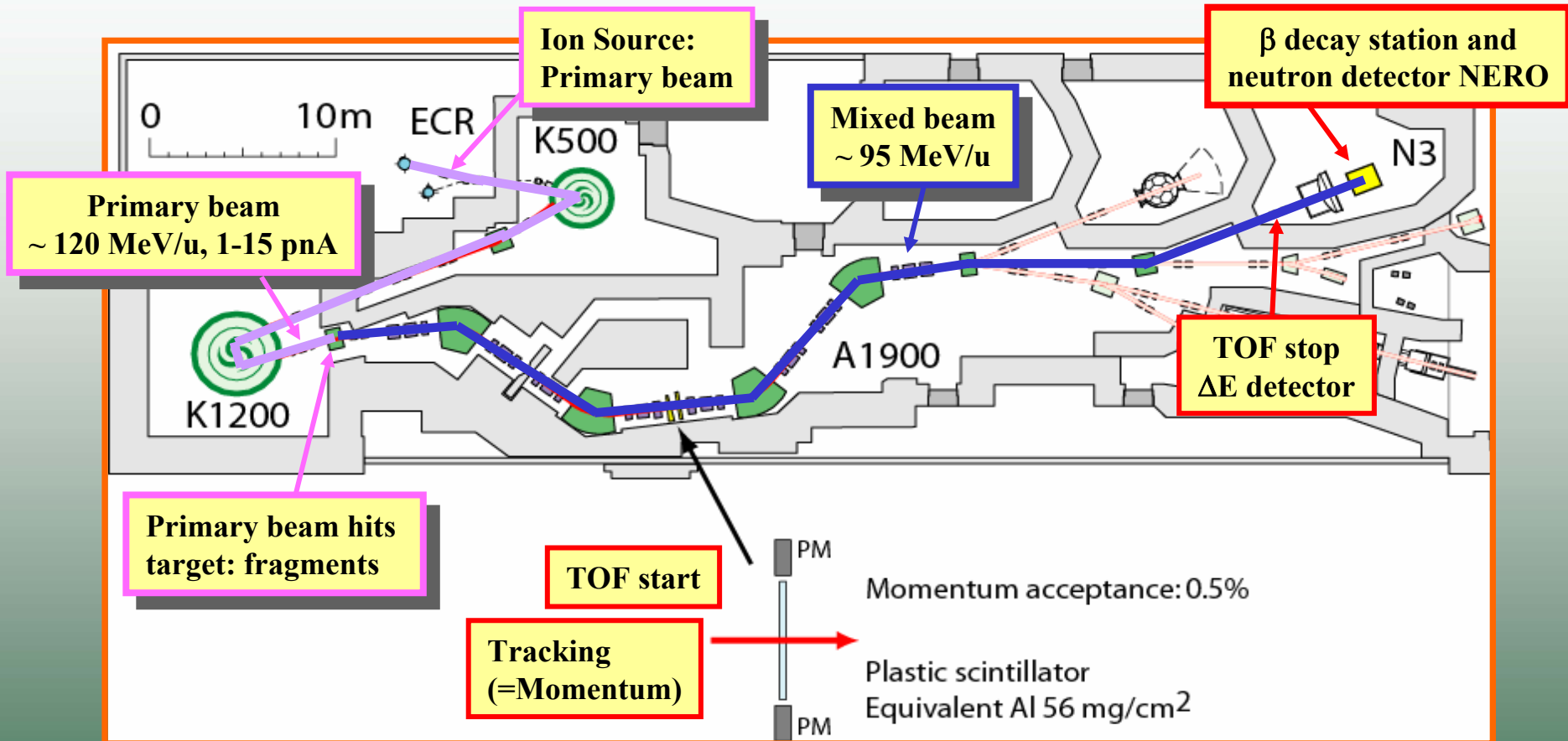
Masses (Sn)
(location of the path)



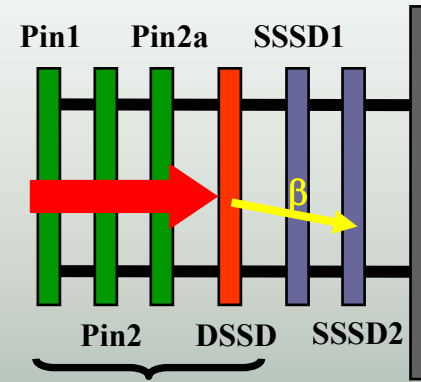
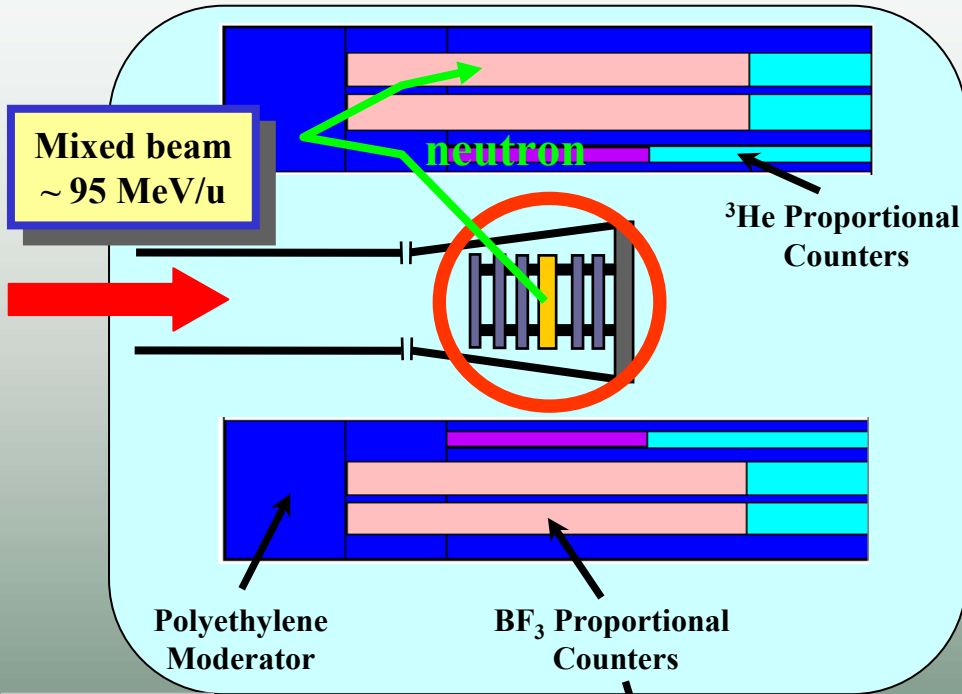
Solar r-process abundances (dots) and abundances using the classical r-process model based on the ETFSI-Q (solid line) mass model and ETFSI-1 (dashed line) mass model



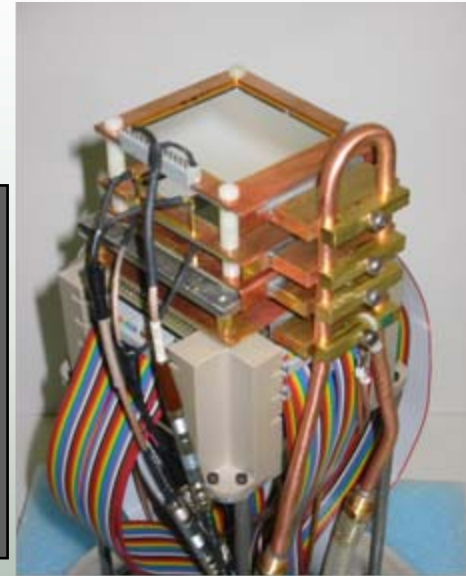
- Only experimental information available for a few isotopes at the N=50 and N=82 shell closure
- So far, we need to rely heavily on theoretical calculations:
 QRPA + Mass models (FRDM, ETFSI)



β decay station and neutron detector NERO



$$\text{TKE} = \text{Pin1} + \text{Pin2} + \text{Pin2a} + \text{DSSD}$$

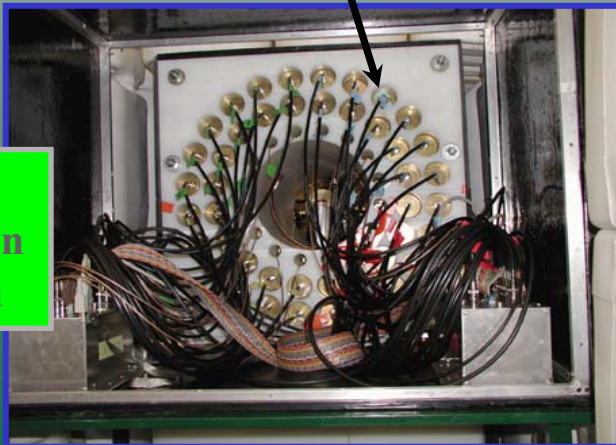


Implant and
 β decay detector

Half-lives
 $T_{1/2}$

Neutron
detector

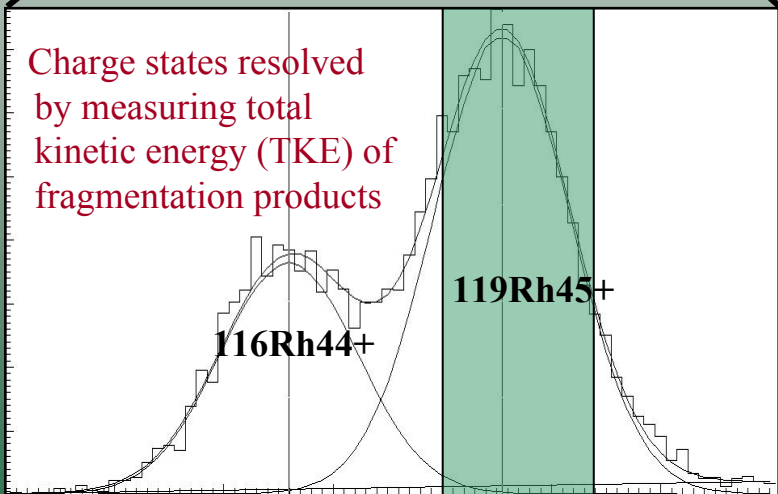
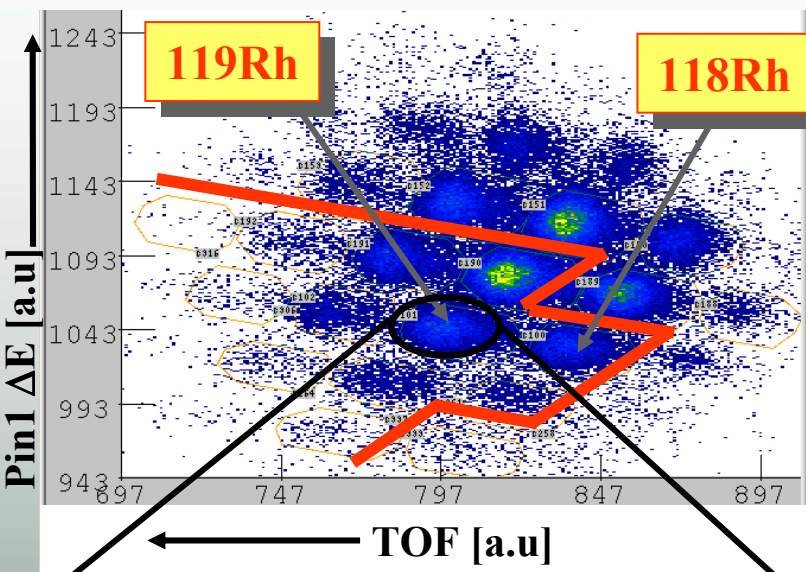
β -delayed
neutron emission
probabilities P_n



Neutron Emission Ratio Observer (NERO):

- 60 counters total (16 ^3He , 44 BF_3)
- 60 cm x 60 cm x 80 cm polyethylene block
- Efficiency ~40%

Particle ID: decays



Sn (50)	Sn120	Sn121	Sn122	Sn123	Sn124	Sn125	Sn126	Sn127	Sn128	Sn129	Sn130	Sn131	Sn132	Sn133
In (49)	In119	In120	In121	In122	In123	In124	In125	In126	In127	In128	In129	In130	In131	In132
Cd (48)	Cd118	Cd119	Cd120	Cd121	Cd122	Cd123	Cd124	Cd125	Cd126	Cd127	Cd128	Cd129	Cd130	Cd131
Ag (47)	Ag117	Ag118	Ag119	Ag120	Ag121	Ag122	Ag123	Ag124	Ag125	Ag126	Ag127	Ag128	Ag129	Ag130
Pd (46)	Pd116	Pd117	Pd118	Pd119	Pd120	Pd121	Pd122	Pd123	Pd124	Pd125	Pd126	Pd127	Pd128	Pd129
Rh (45)	Rh115	Rh116	Rh117	Rh118	Rh119	Rh120	Rh121	Rh122	Rh123	Rh124	Rh125	Rh126	Rh127	Rh128
Ru (44)	Ru114	Ru115	Ru116	Ru117	Ru118	Ru119	Ru120	Ru121	Ru122	Ru123	Ru124	Ru125	Ru126	Ru127
Tc (43)	Tc113	Tc114	Tc115	Tc116	Tc117	Tc118	Tc119	Tc120	Tc121	Tc122	Tc123	Tc124	Tc125	Tc126
Mo (42)	Mo112	Mo113	Mo114	Mo115	Mo116	Mo117	Mo118	Mo119	Mo120	Mo121	Mo122	Mo123	Mo124	Mo125
Nb (41)	Nb111	Nb112	Nb113	Nb114	Nb115	Nb116	Nb117	Nb118	Nb119	Nb120	Nb121	Nb122	Nb123	Nb124
Zr (40)	Zr110	Zr111	Zr112	Zr113	Zr114	Zr115	Zr116	Zr117	Zr118	Zr119	Zr120	Zr121	Zr122	Zr123

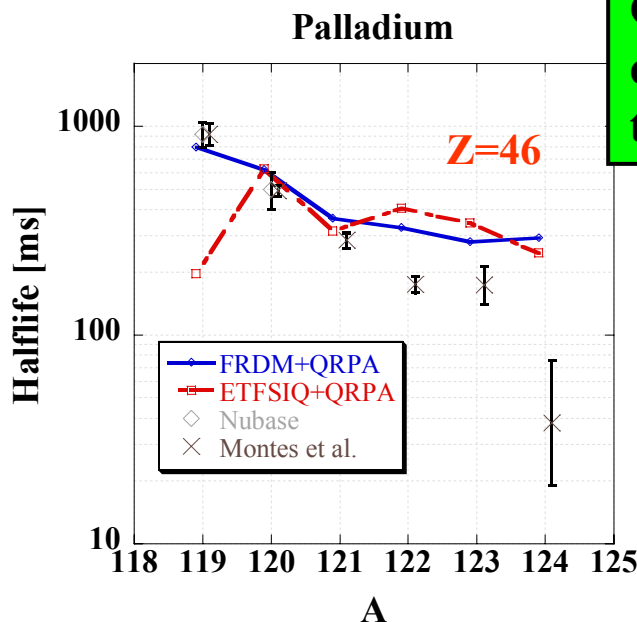
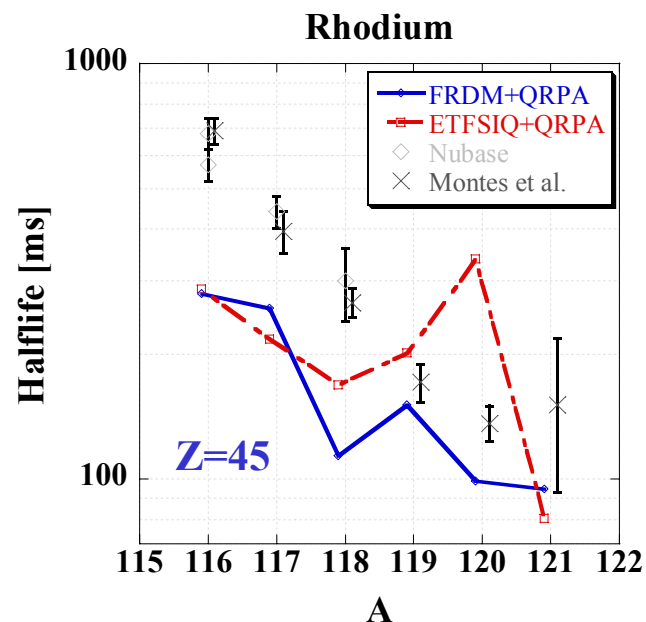
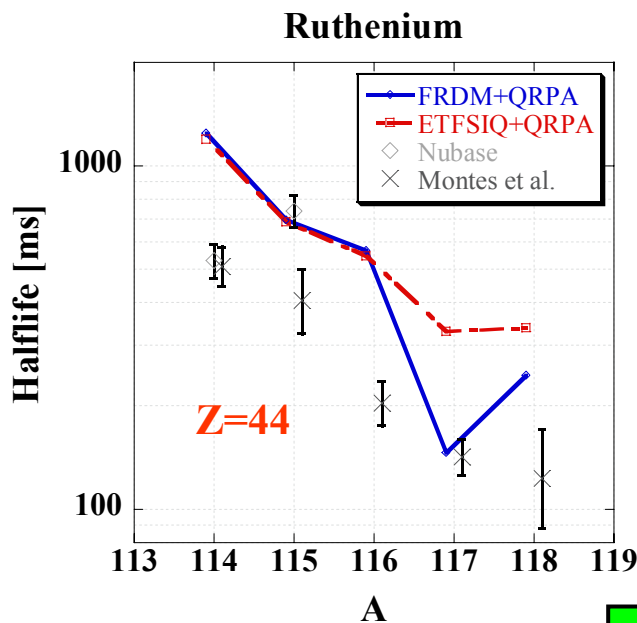
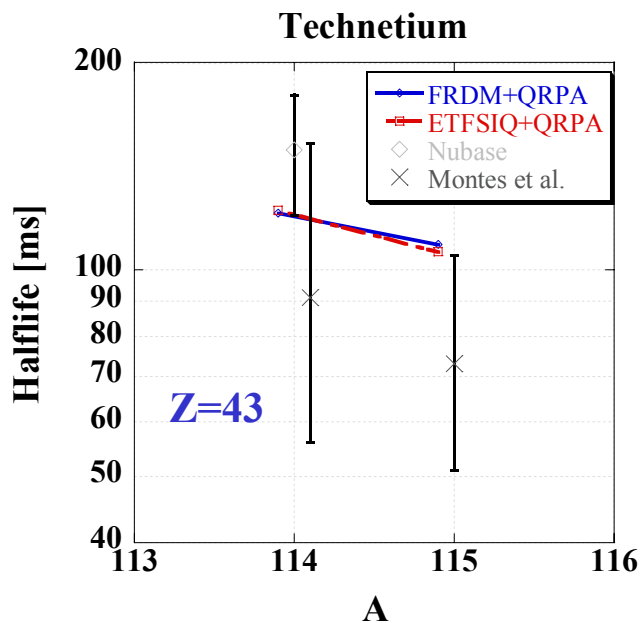
70 71 72 73 74 75 76 77 78 79 80 81 82 83

r-process

- New $T_{1/2}$**
- Improved $T_{1/2}$**
- New P_n**

TKE [a.u.]

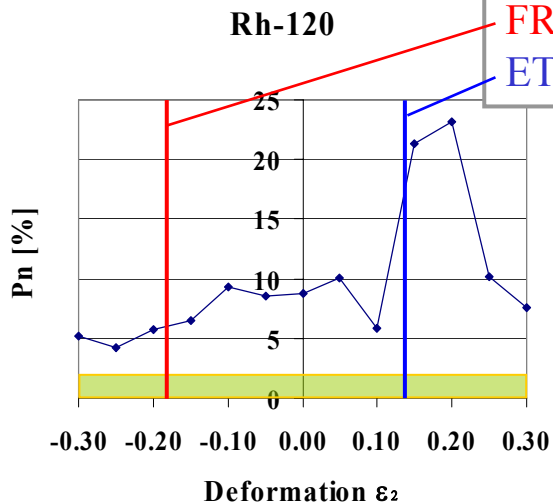
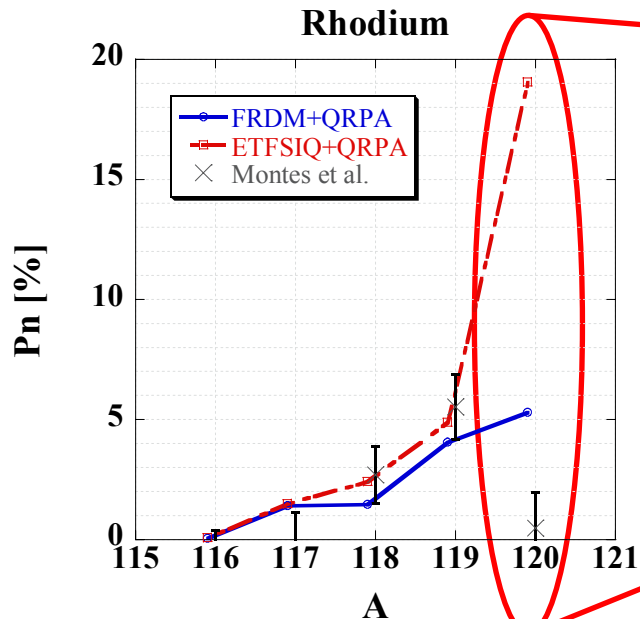




Our $T_{1/2}$ contribution

Good agreement between experimental and theoretical values

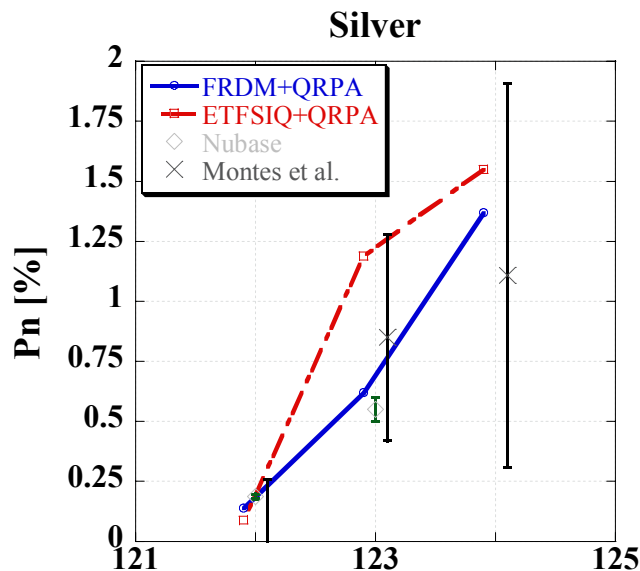




Montes, Pfeiffer, Kratz
Mainz/JINA

Good agreement between experimental and theoretical values

Only surprise ^{120}Rh

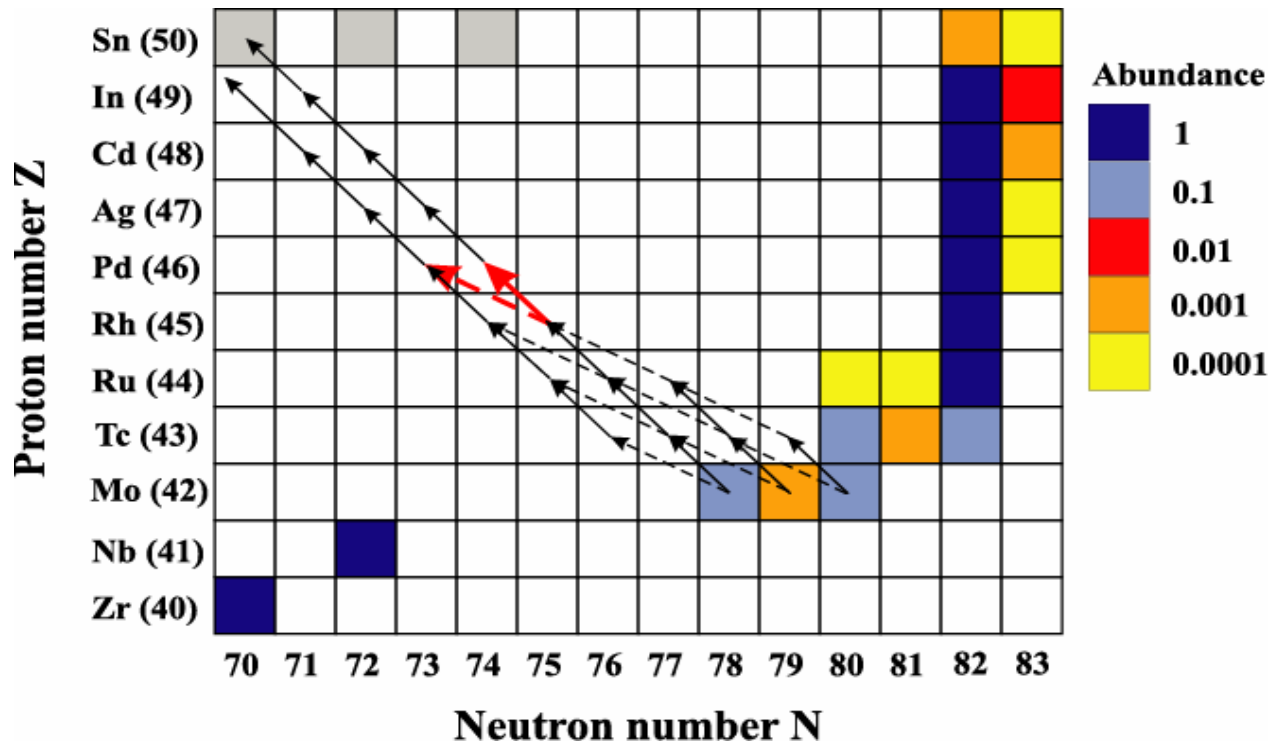


$T_{1/2}$ and Pn rough indicators of nuclear structure ...

Montes et al.

**Deformation and Q value alone are not enough...
s.p. energy levels?
Sn values?**

Calculated isotopic abundances just before freezeout



Branchings modify final abundance before freezeout

Abundance ratio $^{120}\text{Sn}/^{119}\text{Sn}$ is increased by 60% but.....

- Nuclear physics is needed to make full use of astronomical observations and to experimentally constrain r-process models
- Experimental nuclear data (and good theoretical models) are needed
- Good agreement between predictions and experimental values (less than a factor of 3) in this mass region
- Except $^{120}\text{Rh } P_n$ value: not explained by quadrupole deformation or mass uncertainty. Incorrect placement of the dominant GT feeding ($\nu g_{7/2} \rightarrow \pi g_{9/2}$)
- Effect in the r-process? Isotopic ratio $^{120}\text{Sn}/^{119}\text{Sn}$



Collaboration

Michigan State University

Fernando Montes

Alfredo Estrade

Paul Hosmer

Sean Liddick

Paul Mantica

Colin Morton

W. F. Mueller

Michelle Ouellette

Eric Pellegrini

Peter Santi

Hendrik Schatz

Andreas Stolz

Bryan Tomlin

Univ. of Maryland/ANL

W.B. Walters

Univ. of Mainz

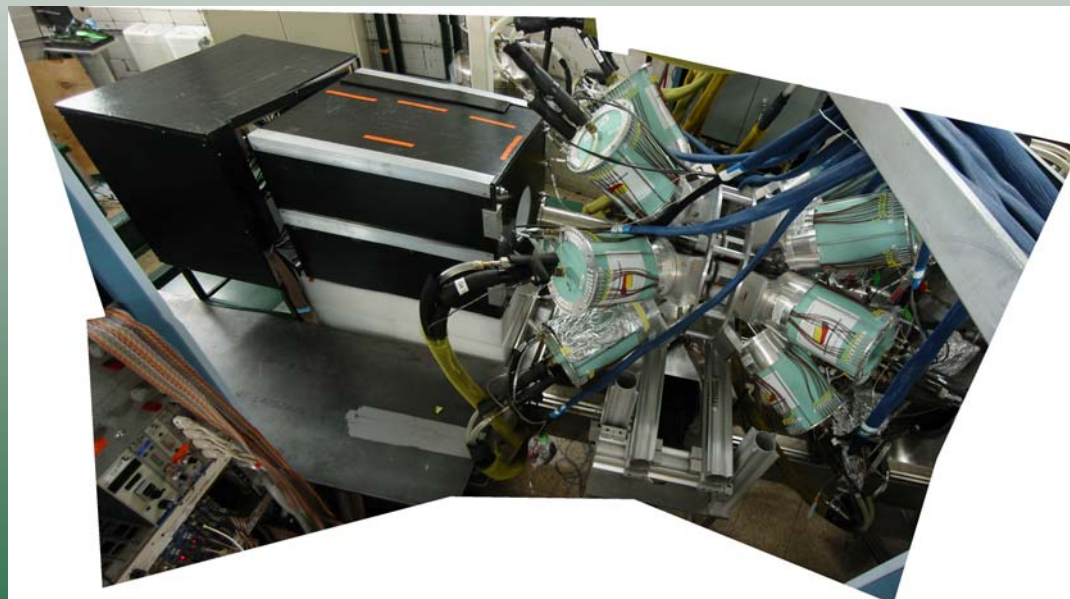
O. Arndt, K.-L. Kratz, B. Pfeiffer

Pacific Northwest Laboratory

P. Reeder

Notre Dame

A. Aprahamian, A. Woehr



**Over-prediction of the experimental $T_{1/2}$ for $N \geq 74$ Pd isotopes...
Uncertainty in the input parameters???**

1. ε_2 does not explain it..
2. An increase in the Q_β value better reproduce the experimental results..

- Large Q_β value ^{130}Cd decay best reproduced by mass models with N=82 shell quenching (Dillmann et al. PRL 91 (2003) 162503)
- Systematic Variation of $E(2^+)$ for ^{120}Pd : hints to the absence of shell quenching (Walters et al. Phys. Rev. C (2004) 034314)

$T_{1/2}$ and Pn rough indicators of nuclear structure ...

Increase in Q_β consistent with a weakening of the neutron shell-closure seen by $N \geq 74$ Pd but.....

